

# Measurement of the extragalactic background light imprint on the spectra of the brightest blazars observed with H.E.S.S.★

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Received 10 September 2012 / Accepted 30 November 2012

## ABSTRACT

The extragalactic background light (EBL) is the diffuse radiation with the second highest energy density in the Universe after the cosmic microwave background. The aim of this study is the measurement of the imprint of the EBL opacity to  $\gamma$ -rays on the spectra of the brightest extragalactic sources detected with the High Energy Stereoscopic System (H.E.S.S.). The originality of the method lies in the joint fit of the EBL optical depth and of the intrinsic spectra of the sources, assuming intrinsic smoothness. Analysis of a total of  $\sim 10^5$   $\gamma$ -ray events enables the detection of an EBL signature at the  $8.8\sigma$  level and constitutes the first measurement of the EBL optical depth using very-high energy ( $E > 100$  GeV)  $\gamma$ -rays. The EBL flux density is constrained over almost two decades of wavelengths [0.30  $\mu\text{m}$ , 17  $\mu\text{m}$ ] and the peak value at 1.4  $\mu\text{m}$  is derived as  $\lambda F_{\lambda} = 15 \pm 2_{\text{stat}} \pm 3_{\text{sys}}$  nW m<sup>-2</sup> sr<sup>-1</sup>.

**Key words.** gamma rays: galaxies – cosmic background radiation – BL Lacertae objects: general

## 1. Introduction

The extragalactic background light (EBL) is the second most intense diffuse radiation in the Universe, and its spectral energy

distribution is composed of two bumps: the cosmic optical background (COB) and the cosmic infra-red background (CIB). The former is mainly due to the radiation emitted by stellar nucleosynthesis in the optical (O) to near infrared (IR), while the latter stems from UV-optical light absorbed and re-radiated by dust in the IR domain (for a review, see Hauser & Dwek 2001).

Direct measurements of the EBL flux density prove to be difficult, mainly because foreground contamination, e.g. by the zodiacal light, can result in an overestimation. Strict lower limits

\* Appendix A is available in electronic form at <http://www.aanda.org>

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have been derived from integrated galaxy counts (see, e.g., Madau & Pozzetti 2000; Fazio et al. 2004; Dole et al. 2006, for more details). The limits derived from direct measurement in the near IR domain typically are one order of magnitude above the lower limits from source counts.

Strong constraints on the EBL density are derived using extragalactic very high energy (VHE,  $E > 100$  GeV)  $\gamma$ -ray sources. VHE  $\gamma$ -rays interact with O-IR photons via electron-positron pair production, resulting in an attenuated flux that is detected on Earth (Nikishov 1962; Jelley 1966; Gould & Schröder 1967). Assuming that there is no intrinsic break in the energy range of interest (as in Stecker et al. 1992) and that the hardness of the spectrum is limited, stringent upper limits on the EBL opacity to  $\gamma$ -rays have been derived (e.g. Aharonian et al. 2006c; Mazin & Raue 2007). Studies exploiting *Fermi*-LAT measurements as templates of the intrinsic spectra have also recently been performed (Georganopoulos et al. 2010; Orr et al. 2011; Meyer et al. 2012). Current models of the EBL are in close agreement with these limits, and they converge on a peak value of the stellar component  $\lambda F_{\lambda} \sim 12$  nW m<sup>-2</sup> sr<sup>-1</sup>, yielding a consistent value for the opacity to  $\gamma$ -rays (see, e.g., Domínguez et al. 2011).

The attenuation by the EBL is expected to leave a unique, redshift dependent and energy dependent imprint on the VHE spectra. While at energies above  $E \gtrsim 5$  to 10 TeV (depending on the redshift of the source) a sharp cut-off is expected resulting from the CIB, a weaker modulation should imprint the spectra in the energy range between  $\sim 100$  GeV and  $\sim 5$ –10 TeV, resulting from the rise and fall of the first peak of the EBL, the COB (Aharonian et al. 1999, 2003). A significant detection of this modulation, localized in a relatively narrow energy range, requires studying high-quality spectra, as, e.g., measured during the strong flux outburst of PKS 2155-304 in 2006 (Aharonian et al. 2007c), under the assumption that the intrinsic spectra are smooth over the energy range being studied.

This signature is searched for in the spectra of the brightest extragalactic blazars detected by H.E.S.S. with a maximum likelihood method, leaving the parameters of the intrinsic spectra free. The originality and the strength of the technique lie in the joint fit of the EBL optical depth and of the intrinsic spectra of the sources, fully accounting for intrinsic curvature. This derivation of the EBL optical depth with H.E.S.S. data does not rely on constraints on the intrinsic spectrum from assumptions about the acceleration mechanism and results in a *measurement* of the optical depth, compared to the upper limits derived in previous studies.

The sample of blazars studied in this paper, the data analysis, and the spectral fitting method are described in Sect. 2. In Sect. 3, the results are presented and the systematic uncertainties are discussed. Finally, the results of this analysis are compared with the current constraints in Sect. 4.

## 2. Analysis of H.E.S.S. data

### 2.1. Reduction of H.E.S.S. data

The high energy stereoscopic system (H.E.S.S.) is an array of four imaging atmospheric Cherenkov telescopes located 1800 m above sea level, in the Khomas Highland, Namibia (23°16'18''S, 16°30'01''E). The Cherenkov light emitted by VHE-particle-induced showers in the atmosphere is focussed with 13 m diameter optical reflectors onto ultra fast cameras (Bernlöhr et al. 2003; Hinton 2004). Each camera consists of 960 photomultipliers equipped with Winston cones to maximize the collection of light. The coincident detection of a shower with at least two

telescopes improves the  $\gamma$ /hadron separation (Funk et al. 2004; Aharonian et al. 2006a).

The data sets studied in this paper were selected with standard quality criteria (weather and stability of the instruments as in Aharonian et al. 2006a), and the main analysis was performed with *Model analysis* (de Naurois & Rolland 2009). Based on a maximum likelihood method that compares the recorded images with simulated  $\gamma$ -rays, this analysis improves the  $\gamma$ /hadron separation with respect to the standard Hillas analysis method (see e.g. Aharonian et al. 2006a), especially for low energies.

The lowest photo-electron threshold of 40 p.e. per camera image after cleaning (Loose Cuts, de Naurois & Rolland 2009) was adopted to cover the largest possible energy range. The on-events were taken from circular regions around the sources with a radius of 0.11°. The background was estimated with the conventional reflected regions method (Aharonian et al. 2006a). A minimum of three operating telescopes was required to derive the spectrum of a source, the redundancy allowing an improved reconstruction of the direction and energy of the  $\gamma$ -rays.

A cross-check was performed with the standard multi-variate analysis (MVA) described in Ohm et al. (2009) and an independent calibration, yielding consistent results.

### 2.2. Sample of sources

The detection of a subtle absorption feature, such as the effect of the EBL, relies on spectra measured with great accuracy, motivating the study of the extensively observed, bright H.E.S.S. blazars. A cut on the detection significance (Li & Ma 1983) of  $10\sigma$  yielded a sample of seven blazars: Mrk 421, PKS 2005-489, PKS 2155-304, 1ES 0229+200, H 2356-309, 1ES 1101-232, and 1ES 0347-121.

Mrk 421 is the first extragalactic source ever detected in the VHE energy domain (Punch et al. 1992). This highly variable BL Lac object is observed by H.E.S.S. at large zenith angles (Aharonian et al. 2005a), yielding a high energy threshold around 1 TeV but also, with a large effective area at higher energies, photons up to  $\sim 40$  TeV. Thus, even considering the low redshift of the source,  $z = 0.031$  (Ulrich et al. 1975), the EBL significantly impacts its observed spectrum, with an optical depth  $\tau(E = 10$  TeV,  $z = 0.031) \sim 1$ .

PKS 2005-489 ( $z = 0.071$ , Falomo et al. 1987) and H 2356-309 ( $z = 0.165$ , Jones et al. 2009) are two blazars at the  $\sim 2\%$  of the Crab nebula flux level, detected by H.E.S.S. since it went into operation (Aharonian et al. 2006b, 2005b). While the latter does not show any sign of spectral variability (Abramowski et al. 2010a), an intensive observation campaign on the former revealed significant variations (Acero et al. 2010; Abramowski et al. 2011).

Together with H 2356-309, 1ES 1101-232 has already been used for EBL studies. With a measured photon index smaller than three for a redshift of 0.186 (Remillard et al. 1989), the spectrum of this source largely contributed to the stringency of the upper-limit derived by Aharonian et al. (2006c). A dedicated study published in 2007 did not reveal any significant flux variations over the observation period between 2004 and 2005 (Aharonian et al. 2007d).

PKS 2155-304 ( $z = 0.116$ , Falomo et al. 1993) is the brightest extragalactic source in the Southern sky, and it has been widely studied with H.E.S.S. (Aharonian et al. 2005c,d, 2007c, 2009b,a; Abramowski et al. 2010b, 2012). It exhibited a spectacular flux outburst in July 2006 (Aharonian et al. 2007c), with a flux so high that the number of detected  $\gamma$ -rays exceeds by far the cumulated excess from all the other H.E.S.S. extragalactic

**Table 1.** Data sets on VHE blazars detected by H.E.S.S. that are used for this study of the EBL.

Data set	$N_\gamma$	$\sigma$	$E_{\min}-E_{\max}$ [TeV]
Mrk 421 (1)	3381	96.7	0.95–41
Mrk 421 (2)	5548	135	0.95–37
Mrk 421 (3)	5156	134	0.95–45
PKS 2005-489 (1)	1540	25.3	0.16–37
PKS 2005-489 (2)	910	28.9	0.18–25
PKS 2155-304 (2008)	5279	99.2	0.13–19
PKS 2155-304 (1)	3499	93.0	0.13–5.7
PKS 2155-304 (2)	3470	116	0.13–9.3
PKS 2155-304 (3)	9555	186	0.13–14
PKS 2155-304 (4)	4606	132	0.18–4.6
PKS 2155-304 (5)	11 901	219	0.13–5.7
PKS 2155-304 (6)	6494	166	0.15–5.7
PKS 2155-304 (7)	8253	191	0.20–7.6
1ES 0229+200	670	12.6	0.29–25
H 2356-309	1642	21.2	0.11–34
1ES 1101-232	1268	17.8	0.12–23
1ES 0347-121	604	13.5	0.13–11

**Notes.** For highly variable sources, the data are divided into smaller subsets that are indexed in Col. 1 and correspond to restricted flux ranges. The photon excess, detection significance, and energy range of the spectra (in TeV) are given in Cols. 2–4, respectively.

sources. This study focusses on the high statistics data set from July 2006 and from a multi-wavelength campaign performed in 2008, where the low state of the source was measured with high precision (Aharonian et al. 2009b). These detailed high quality spectra of PKS 2155-304 have not been used to set limits on the EBL so far and are responsible for the most stringent constraints derived in this paper.

1ES 0229+200 and 1ES 0347-121 are characterized by their redshift of 0.14 (Schachter et al. 1993) and 0.188 (Woo et al. 2005), respectively. The spectra of these sources (Aharonian et al. 2007a,b) confirmed the EBL limits set by Aharonian et al. (2006c), and their light curves were compatible with constant flux.

For each source, the redshift, excess, significance, and energy range of the detected  $\gamma$ -rays are shown in the first columns of Table 1. Blazars sometimes exhibit spectral changes correlated with the flux (e.g. Abramowski et al. 2010b) that could result in a scatter of the absorption feature estimates. Spectral variations can be particularly important compared to statistical fluctuations for highly significant ( $\geq 30\sigma$ ) sources. To minimize this effect, the data from PKS 2155-304 (high state), Mrk 421, and PKS 2005-489 were divided into several bins in flux with roughly the same logarithmic width and a similar number of  $\gamma$ -rays, using data slices of 28 min duration (runs). This resulted in 7, 3, and 2 bins for the sources, respectively, which are ordered by increasing level of flux and are listed in brackets in Table 1. The observational conditions for the various data sets on a single source vary from one set to another, implying different energy ranges.

### 2.3. Spectral analysis

The spectral analysis of the data sets described in Sect. 2.2 was performed taking the EBL absorption  $e^{-\tau(E,z,n)}$  into account, where the optical depth depends on the EBL density  $n$  and on the energy  $E$  of the  $\gamma$ -rays, emitted by a source located at a redshift  $z$ . The EBL optical depth was scaled with a normalization factor

$\alpha$ , as in Abdo et al. (2010), yielding a spectral model for each source:

$$\phi_z(E) = \phi_{\text{int}}^\alpha(E) \times \exp(-\alpha \times \tau(E, z, n)) \quad (1)$$

where  $\phi_{\text{int}}^\alpha(E)$  is the intrinsic spectrum of the source, i.e. the de-absorbed spectrum assuming an EBL optical depth scaled by  $\alpha$ .

The template chosen for the EBL density  $n$  is the model of Franceschini et al. (2008), hereafter FR08, which is representative of the current state of the art of EBL modelling and for which the optical depth is finely discretized in energy and redshift<sup>1</sup>. The EBL normalization factor  $\alpha$ , defined in Eq. (1), is thus an estimator of the ratio between the measured and template opacities  $\tau_{\text{measured}}/\tau_{\text{FR08}}$ . The particular choice of optical depth modelling only has a minor impact on the reconstruction of the EBL flux density, and the systematic uncertainty resulting from this choice is estimated in Sect. 3.3.

The functional form of the intrinsic spectrum  $\phi_{\text{int}}^\alpha(E)$  assumed in this study is taken from very general considerations about the source physics. Blazars spectral energy distributions are indeed commonly described with a leptonic emission, e.g. with synchrotron self Compton models (Band & Grindlay 1985). In the VHE range, a smooth and concave spectrum is expected with the possible addition of a cut-off arising from the Klein-Nishina effect or a cut-off in the underlying electron distribution. The concurrent hadronic scenarios result in a smooth spectrum (see e.g. Mannheim 1993; Beall & Bednarek 1999; Aharonian 2000) that closely resembles the leptonic spectra in the VHE range. At the first order, the intrinsic spectra are described with the most natural functional form for a non-thermal emission: a power law (PWL), i.e. a linear function in log-log scale. To test for the presence of intrinsic curvature, the next order of complexity is readily achieved using the log-parabola (LP), which is the equivalent of the parabola in log-log scale. The exponential cut-off hypothesis is also tested (EPWL), since expected on theoretical grounds, and since the order of the equivalent log-log polynomial would be too high, unreasonably widening the parameter space. The next order of complexity is simply achieved by generalizing the last two models, adding a cut-off to the LP (ELP), and smoothing ( $\gamma < 1$ ) or sharpening ( $\gamma > 1$ ) the cut-off of the EPWL (SEPWL). The exact choice of the intrinsic models, which are detailed in Table 2, does not strongly affect the EBL measurement described hereafter, as shown in Appendix A.

In the following, deviations from concavity are assumed to arise from the EBL absorption term, which is a reasonable assumption as long as the scenarios concurrent to the leptonic emission do not mimic the energy and redshift dependence of the EBL optical depth (but see also Reimer 2007, for other probes such as flat spectrum radio quasars). The energy dependence of the EBL absorption deviates from mere concavity (e.g. Raue & Mazin 2010), and inflection points in the observed spectra, which depend on the redshift of the source, constitute the key imprint that is reconstructed in this study.

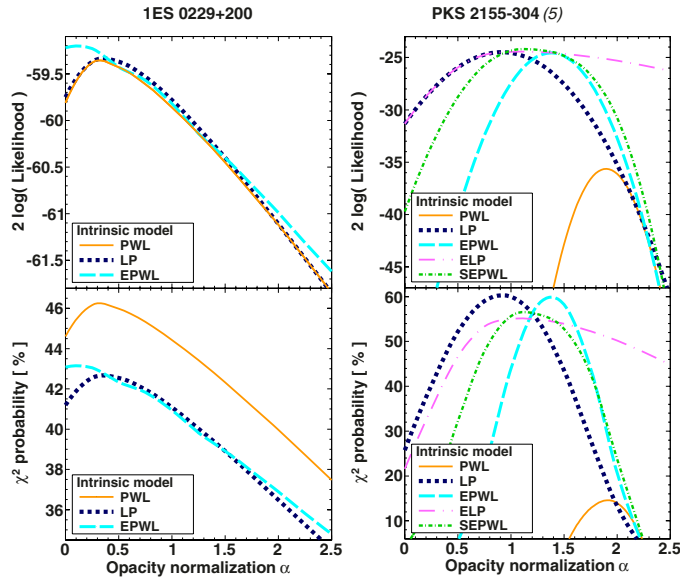
To quantify the amplitude of the EBL signature on H.E.S.S. spectra, the maximum likelihood method developed by Abdo et al. (2010) was adapted. Likelihood profiles were computed as a function of  $\alpha$  for each data set and each smooth intrinsic spectral model given in Table 2, with  $\alpha$  ranging from 0 to 2.5. For each value of  $\alpha$ , the models  $\phi_z(E)$  were fitted to the data

<sup>1</sup> The optical depth derived by FR08 is tabulated from  $z_0 = 10^{-3}$  to  $z_1 = 2$  in steps of  $\delta z = 10^{-3}$  and from  $E_0 = 20$  GeV to  $E_1 = 170$  TeV for 50 logarithmic steps. An interpolation in energy is performed for the spectral analysis.

**Table 2.** Smooth functions describing the intrinsic spectra of the sources studied in this paper.

Name	Abbrev.	Function <sup>2</sup>
Power law	PWL	$\phi_0(E/E_0)^{-\Gamma}$
Log parabola	LP	$\phi_0(E/E_0)^{-a-b\log(E/E_0)}$
Exponential cut-off power law	EPWL	$\phi_0(E/E_0)^{-\Gamma} \exp(-E/E_{\text{cut}})$
Exponential cut-off log parabola	ELP	$\phi_0(E/E_0)^{-a-b\log(E/E_0)} \exp(-E/E_{\text{cut}})$
Super exponential cut-off power law	SEPWL	$\phi_0(E/E_0)^{-\Gamma} \exp(-(E/E_{\text{cut}})^\gamma)$

**Notes.** <sup>(2)</sup> The reference energy  $E_0$  is set to the decorrelation energy of the spectrum.



**Fig. 1.** *Top panels:* likelihood profiles, as a function of the optical depth normalization for the different intrinsic models detailed in the legend. The examples of the data sets on 1ES 0229+200 and PKS 2155-304 (5) are shown on the left and right, respectively. *Bottom panels:* corresponding  $\chi^2$  probabilities as a function of the optical depth normalization. The PWL and SEPWL models are the spectral models chosen to describe the spectra of 1ES 0229+200 and the fifth data set on PKS 2155-304, respectively.

with the intrinsic spectral parameters free in the minimization procedure. The best fit maximum likelihood  $\mathcal{L}$  was converted<sup>3</sup> into an equivalent  $\chi^2 = -2 \log \mathcal{L}$  allowing the goodness of the fit to be assessed with the conventional  $\chi^2$  probability as a function of  $\alpha$ .

An unconventional procedure was set up to select a model. It ensures that the intrinsic curvature is fully taken into account and that the extrinsic curvature due to the EBL absorption is not over-estimated. Normally, the model with fewest parameters would be used unless a model with one extra parameter is statistically preferred. Here, the model with the highest  $\chi^2$  probability was selected, regardless of the value of  $\alpha$  for which this maximum is reached. Cases where two models had comparable maximum  $\chi^2$  probabilities are discussed in the following.

Figure 1 shows the likelihood profiles and the  $\chi^2$  probability profiles derived with the smooth intrinsic spectral models given in Table 2 for the data sets on 1ES 0229+200 and

**Table 3.** Spectral modelling of the data sets used to derive the likelihood profiles.

Data set	Spectral model	$\chi^2(\alpha_0)$ / d.o.f.
Mrk 421 (1)	ELP	21.5 / 31
Mrk 421 (2)	ELP	46.8 / 30
Mrk 421 (3)	ELP	34.8 / 28
PKS 2005-489 (1)	LP	49.5 / 60
PKS 2005-489 (2)	LP	31.8 / 46
PKS 2155-304 (2008)	ELP	21.9 / 37
PKS 2155-304 (1)	PWL	32.3 / 31
PKS 2155-304 (2)	SEPWL	25.3 / 28
PKS 2155-304 (3)	SEPWL	35.2 / 31
PKS 2155-304 (4)	SEPWL	19.1 / 21
PKS 2155-304 (5)	SEPWL	24.3 / 27
PKS 2155-304 (6)	LP	29.2 / 21
PKS 2155-304 (7)	SEPWL	13.6 / 13
1ES 0229+200	PWL	60.1 / 60
H 2356-309	LP	70.2 / 61
1ES 1101-232	PWL	62.6 / 69
1ES 0347-121	ELP	31.7 / 35

**Notes.** The spectral models (see Sect. 2.3) are given in Col. 2, where the acronyms PWL, LP, EPWL, ELP, and SEPWL are explained in Table 2. The  $\chi^2$  for the best fit EBL optical depth normalization  $\alpha_0$  and the number of degrees of freedom d.o.f. are given in Col. 3.

PKS 2155-304 (5). In the first case, the likelihood profile derived with the PWL (two parameters) model does not significantly differ from those obtained with a LP (three parameters) or an EPWL (three parameters), but the decrease in the number of degrees of freedom with increasing complexity favours the PWL in term of  $\chi^2$  probability, as shown in the bottom panel. In this case, the conventional method and our approach select the same model.

In the second case, corresponding to the fifth data set on PKS 2155-304, the LP and the EPWL significantly improve the fit compared to the PWL, in terms of maximum likelihood and of maximum  $\chi^2$  probability. The LP profile and the EPWL profile have a similar maximum likelihood and, equivalently, (since the two models have the same number of free parameters) a similar maximum  $\chi^2$  probability. Instead of performing an arbitrary choice between the LP and the EPWL, the profiles of the SEPWL, which generalizes the EPWL, and of the ELP, which generalizes the LP, were computed. According to our approach, the model with the highest maximum  $\chi^2$  probability, in this case the SEPWL, was then selected.

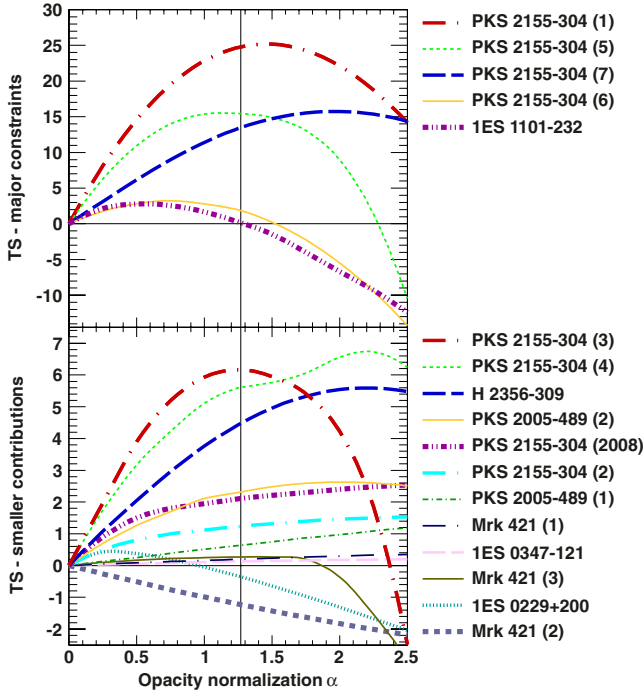
As shown in Fig. A.2, using the common criteria with a significance level of  $2\sigma$  yields results in agreement with the unconventional method described here. The drawback of our approach is a less significant measurement with regards to the conventional method, due to the widening of the studied parameter space and the consequently larger statistical uncertainties. The intrinsic spectral models that were selected for each data set are given in Table 3. The impact of the selection of the intrinsic model on the final result is investigated in Fig. A.2 and is included in the systematics (Sect. 3.3).

### 3. Results

#### 3.1. Measurement of the EBL optical depth normalization

For each data set, the likelihood  $\mathcal{L}(\alpha)$  of the EBL optical depth normalization and of the intrinsic spectral model is compared to the hypothesis of a null EBL absorption  $\mathcal{L}(\alpha = 0)$ . The

<sup>3</sup> The log-likelihood is set to zero if the measured number of events matches the expected number of events in each bin.



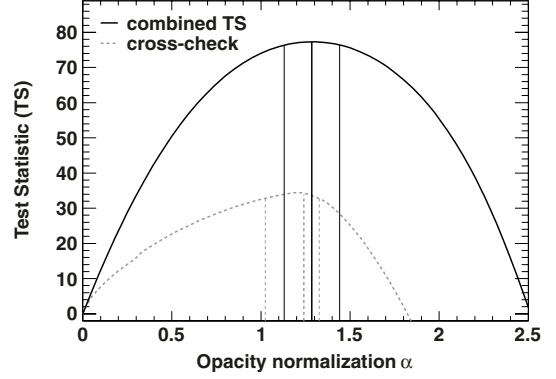
**Fig. 2.** Test statistic as a function of the normalized EBL optical depth for the intrinsic spectral models described in Table 3. The TS profiles are sorted by contribution to the measurement and the top panel shows the most constraining data sets, while the bottom panel shows the less constraining contributions. The vertical line indicates the best fit value derived in this study. Note the different scale on the vertical axis in the upper and the lower panel.

test statistic (TS), defined by the likelihood ratio test  $TS = 2 \log [\mathcal{L}(\alpha = \alpha_0) / \mathcal{L}(\alpha = 0)]$ , is shown for each data set in Fig. 2.

With a total  $\gamma$ -ray excess of the order of 50 000 events, PKS 2155-304 makes a major contribution to the EBL measurement. A maximum TS superior to 16 is achieved for the data sets (1), (5), and (7), meaning that a null EBL optical depth is rejected at the  $4\sigma$  level by each of these data sets. An EBL optical depth scaled up by a factor two is rejected at the  $3\sigma$  level by both the data set (6) on PKS 2155-304 and the one on 1ES 1101-232. This constraint is not surprising since 1ES 1101-232 was already the most constraining source used by Aharonian et al. (2006c) to derive an upper-limit on the EBL opacity. The bottom panel of Fig. 2 shows the less constraining contributions. Though less significant individually, their combination contributes to roughly a third of the total TS and enables a null EBL optical depth to be rejected at the  $\sim 5\sigma$  level.

The total TS shown in Fig. 3, i.e. the sum of the individual ones presented in Fig. 2, is maximum for  $\alpha_0 = 1.27^{+0.18}_{-0.15}$ , at  $\sqrt{\Delta TS} \sim 1.8\sigma$  above the unscaled FR08 template. The upper and lower standard deviations correspond to a variation of  $\Delta TS = 1$  from the maximum test statistic  $TS = 77.3$ . The EBL optical depth template scaled up by a factor  $\alpha_0$  is preferred at the  $8.8\sigma$  level to a null optical depth, where the Gaussian significance is approximated by the square root of the likelihood ratio test.

No outlier is present in the set of individual profiles, with best fit values of  $1.44 \pm 0.29$  ( $0.6\sigma$ ),  $1.23 \pm 0.34$  ( $0.1\sigma$ ),  $1.97 \pm 0.48$  ( $1.5\sigma$ ),  $0.75 \pm 0.42$  ( $1.2\sigma$ ), and  $0.48 \pm 0.29$  ( $1.6\sigma$ ) for the five most constraining data sets PKS 2155-304 (1, 5, 7, 6) and 1ES 1101-232, respectively, where the number in brackets indicate the deviation to the best fit value  $\alpha_0 = 1.27$ . Similarly, the less constraining contributions do not differ by more than  $\sqrt{TS_{\max} - TS(\alpha_0)} \lesssim 1.5\sigma$  from the best fit value.



**Fig. 3.** Combined test statistic as a function of the normalized EBL optical depth. The results obtained with the *Model analysis* are shown with a black line and the cross-check led with the MVA analysis is shown with the dashed grey line. The best fit value and  $1\sigma$  statistical uncertainties are shown with the vertical lines.

The total TS derived with the MVA analysis and an independent calibration is shown in Fig. 3. Though less significant, with a maximum TS of 33.9, the best fit value  $1.24^{+0.09}_{-0.22}$  is in close agreement with the optical depth normalization derived with the *Model analysis*. The larger energy range covered by the latter (60% wider in logarithmic scale) accounts for the difference in maximum TS.

### 3.2. Redshift dependence

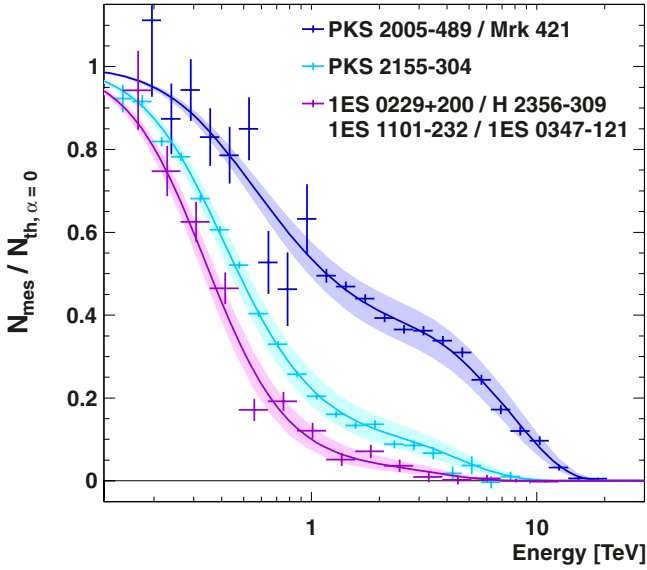
To investigate the redshift dependence of the EBL optical depth normalization, the data set is divided by redshift in three groups. For Mrk 421 and PKS 2005-489, the TS is maximum at  $\alpha(z_1) = 1.6^{+0.5}_{-1.1}$ , for an average redshift of  $z_1 = 0.051$ . The TS of the eight data sets on PKS 2155-304 ( $z_2 = 0.116$ ) peaks at  $\alpha(z_2) = 1.36 \pm 0.17$ . With the four other data sets, a maximum TS is obtained for  $\alpha(z_3) = 0.71^{+0.46}_{-0.29}$ , corresponding to a mean redshift  $z_3 = 0.170$ .

Fitting the decreasing trend of the EBL normalization as a linear function of the redshift yields  $\chi^2/\text{d.o.f.} = 0.41/1$ , which does not significantly improve the fit with regards to a constant fit  $\chi^2/\text{d.o.f.} = 1.83/2$ . A likelihood ratio test prefers the linear fit only at the  $1.1\sigma$  level. Any redshift dependence of the EBL normalization in the redshift range probed is therefore neglected in the following sections.

Given the limited amount of data, the deviations from the best fit EBL normalization  $\alpha_0 = 1.27^{+0.18}_{-0.15}$  can hardly be investigated at the single data set level. For the three above-mentioned groups of sources, the total number of measured events in each energy bin ( $N_{\text{mes}}$ ) is scaled to the expected number of events from the intrinsic spectra ( $N_{\text{th}, \alpha=0}$ ). This ratio is compared in Fig. 4 to the best fit model for the three average redshifts of 0.051, 0.116, and 0.17. Abrupt changes in the amplitude of the statistical uncertainties (e.g. around 1 TeV for the low redshift group: Mrk 421 / PKS 2005-489) are inherent to the grouping of data sets that cover different energy ranges (e.g. the data sets on Mrk 421 start at  $\sim 1$  TeV).

### 3.3. Systematic uncertainty

An extensive investigation was undertaken of the systematic uncertainties arising from the method. Four sources of systematic uncertainties on the EBL optical depth normalization were identified: the analysis chain (background rejection, spectral



**Fig. 4.** Observed number of  $\gamma$ -rays over number of events expected from the intrinsic spectra vs.  $\gamma$ -ray energy. The data sets are grouped by similar redshift and the detected and expected numbers of  $\gamma$ -rays are summed in each energy bin. The best fit EBL absorption is represented by the solid lines for the three redshifts corresponding to the groups of data sets and the shaded areas correspond to the  $\pm 1\sigma$  best fit EBL normalization.

**Table 4.** Sources of systematics and estimated uncertainties on the normalized EBL optical depth  $\alpha_0 = 1.27^{+0.18}_{-0.15 \text{ stat}}$ .

Sources of systematics	Estimated systematics
Analysis chain	0.21
Intrinsic model	0.10
EBL model	0.06
Energy scale	0.05
Total	0.25

**Notes.** A full discussion of the systematic uncertainties can be found in Appendix A.

analysis), the choice of intrinsic models and of the EBL template, as well as the limited knowledge of the energy scale due to the atmosphere. These systematic uncertainties are summarized in Table 4 and detailed in Appendix A.

The total systematic is estimated as  $\sigma_{\text{sys}}(\alpha_0) = 0.25$  and is comparable to the statistical uncertainty on the normalized EBL optical depth  $\alpha_0 = 1.27^{+0.18}_{-0.15 \text{ stat}}$ .

## 4. Discussion

The measurement of the EBL optical depth can be converted to an EBL flux density, but particular attention must be paid to the wavelength range covered.

A  $\gamma$ -ray of energy  $E^*$  and an EBL photon of energy  $\epsilon^*$  tend to produce an electron-positron pair mostly for  $E^*\epsilon^* = (2m_e c^2)^2$  (peak of the cross section, see, e.g., Jauch & Rohrlich 1976). The interaction can occur anywhere along the path of the  $\gamma$ -ray from the source and the relation for the EBL wavelength becomes, in the observer frame,

$$(\lambda_{\text{EBL}}/1 \mu\text{m}) = 1.187 \times (E/1 \text{ TeV}) \times (1+z')^2 \quad (2)$$

with  $z' < z$ , where  $z$  is the redshift of the source and where  $E$  is the  $\gamma$ -ray energy in the observer frame. To derive this relation between the EBL wavelength and the  $\gamma$ -ray energy, the width of the

**Table 5.** EBL wavelength range probed by the data sets used in this study.

Data set	$z$	$E_{\text{min}}-E_{\text{max}}$ [TeV]	$\lambda_{\text{min}}-\lambda_{\text{max}}$ [ $\mu\text{m}$ ]
Mrk 421 (1)	0.031	0.95–41	1.2–49
Mrk 421 (2)	0.031	0.95–37	1.2–44
Mrk 421 (3)	0.031	0.95–45	1.2–53
PKS 2005-489 (1)	0.071	0.16–37	0.22–44
PKS 2005-489 (2)	0.071	0.18–25	0.25–30
PKS 2155-304 (2008)	0.116	0.13–19	0.30–23
PKS 2155-304 (1)	0.116	0.13–5.7	0.19–6.8
PKS 2155-304 (2)	0.116	0.13–9.3	0.19–11
PKS 2155-304 (3)	0.116	0.13–14	0.19–17
PKS 2155-304 (4)	0.116	0.18–4.6	0.19–5.5
PKS 2155-304 (5)	0.116	0.13–5.7	0.27–6.8
PKS 2155-304 (6)	0.116	0.15–5.7	0.19–6.8
PKS 2155-304 (7)	0.116	0.20–7.6	0.22–9.0
1ES 0229+200	0.14	0.29–25	0.45–30
H 2356-309	0.165	0.11–34	0.18–40
1ES 1101-232	0.186	0.12–23	0.20–27
1ES 0347-121	0.188	0.13–11	0.22–13

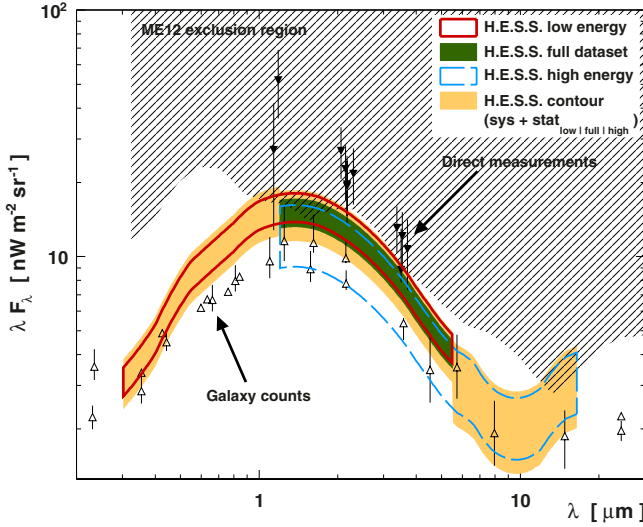
**Notes.** The redshifts of the sources are given in Col. 2. The energy range of the spectra (in TeV) is given in Col. 3, and the EBL wavelengths probed with the subsets are given in Col. 4, where only the peak of the pair-creation cross-section is taken into account.

pair-creation cross-section as a function of energy is neglected. Taking it into account would result in an even wider wavelength coverage for a given  $\gamma$ -ray energy range.

The detection of an EBL flux density scaled up by a factor  $\alpha_0 = 1.27^{+0.18}_{-0.15 \text{ stat}} \pm 0.25_{\text{sys}}$  is then valid in the overlap of the data-set energy ranges  $[(1+z)^2 E_{\text{min}}, E_{\text{max}}]$ , where the factor  $(1+z)^2$  accounts for the redshift dependency in Eq. (2). The measurement that is derived with all data sets is shown by the filled area in Fig. 5 in the wavelength range  $[1.2, 5.5] \mu\text{m}$ , where  $1.2 \mu\text{m}$  (resp.  $5.5 \mu\text{m}$ ) is the counterpart of the low (resp. high) energy bound of the Mrk 421 (resp. PKS 2155-304) data sets, as shown in Table 5.

To probe a wider wavelength range and to ensure the consistency of the modelling below and above  $\sim 1 \mu\text{m}$ , the TSS of data sets with comparable energy ranges were combined. Low EBL-wavelengths between  $0.30$  and  $5.5 \mu\text{m}$  were studied with the combination of the 1ES 0347-121 data set and the six PKS 2155-304 data sets (1, 2, 4, 5, 6, 7) while the large EBL-wavelengths between  $1.2$  and  $17 \mu\text{m}$  were probed by the 1ES 1101-232, 1ES 0229+200, PKS 2005-489, Mrk 421, H 2356-309 data sets, and the two PKS 2155-304 data sets (3, 2008), all described in Table 5. The normalized EBL optical depth measured in the various wavelength ranges and the corresponding EBL flux density are given in Table 6.

The  $1\sigma$  (statistical) contours of the EBL flux density for these two wavelength ranges and for the combination are compared in Fig. 5 to other measurements and limits. The first peak of the EBL flux density, the COB, is entirely constrained by the low and the high energy data sets. The systematic uncertainty is quadratically added to the statistical uncertainty on the measurement with the full data set in the intermediate wavelength range, and to uncertainties on the low and high energy measurements in the extended ranges. The statistical uncertainties remain dominant around  $10 \mu\text{m}$ . In the UV to NIR domain, the systematic uncertainties, which are propagated from the optical depth normalization to the flux density as a single normalization factor, make a non-negligible contribution to the width of the contour.



**Fig. 5.** Flux density of the extragalactic background light versus wavelength. The  $1\sigma$  (statistical) contours derived for several energy ranges are described in the top-right legend. The systematic uncertainty is added quadratically to the statistical one to derive the H.E.S.S. contour. Lower limits based on galaxy counts and direct measurements are respectively shown with empty upward and filled downward pointing triangles (extracted from Gilmore et al. 2012). The region excluded by Meyer et al. (2012) with VHE spectra is represented by the dashed area.

**Table 6.** Measured normalization of the EBL optical depth, corresponding to the  $1\sigma$  (statistical) contours shown in Fig. 5.

$\tau_{\text{measured}}/\tau_{\text{FR08}}$	$\lambda_{\text{min}}-\lambda_{\text{max}}$ $\mu\text{m}$	$\lambda F_{\lambda}(\lambda_{\text{min}})-\lambda F_{\lambda}(\lambda_{\text{max}})$ $\text{nW m}^{-2} \text{sr}^{-1}$
$1.27^{+0.18}_{-0.15}$	1.2–5.5	$14.8^{+2.1}_{-1.7}-4.0^{+0.6}_{-0.5}$
$1.34^{+0.19}_{-0.17}$	0.30–5.5	$3.1 \pm 0.4-4.2^{+0.6}_{-0.5}$
$1.05^{+0.32}_{-0.28}$	1.2–17	$12.2^{+3.7}_{-3.3}-3.2^{+1.0}_{-0.8}$

**Notes.** The second column indicates the wavelength range where this measurement is valid and the third column the corresponding flux densities. The first line corresponds to the full data set. The second and third lines indicate the value derived with smaller data sets focussed on specific energy ranges. The systematic uncertainty on the measurements listed in the first column is 0.25.

The detailed study of the dependence of the systematic uncertainties on the wavelength, based e.g. on deviations from the EBL template model, is beyond the scope of this paper but the comparison of various modellings in a complementary redshift band and wavelength range by The Fermi-LAT Collaboration (Ackermann et al. 2012) supports our choice of template.

The contours lie in between the constraints derived with galaxy counts and the direct measurements. A good agreement with the VHE upper limit derived by Meyer et al. (2012) is also found over the wavelength range covered, with a maximum discrepancy between 1 and 2  $\mu\text{m}$  smaller than the  $1\sigma$  level. The analysis performed enables a measurement of the COB peak flux density of  $\lambda F_{\lambda} = 15.0^{+2.1}_{-1.8} \pm 2.8_{\text{sys}} \text{ nW m}^{-2} \text{sr}^{-1}$  at 1.4  $\mu\text{m}$ , where the peak value and uncertainties are derived by scaling up the FR08 EBL flux density by a factor  $\alpha_0$ . This value is compatible with the previous constraints on the EBL flux density derived with H.E.S.S. data by Aharonian et al. (2006c).

## 5. Summary and conclusion

The spectra of the brightest blazars detected by H.E.S.S. were investigated for an EBL absorption signature. Assuming intrinsic spectral smoothness, the intrinsic spectral curvature was carefully disentangled from the EBL absorption effect. The EBL imprint is detected at an  $8.8\sigma$  level, which constitutes the first measurement of the EBL optical depth using VHE  $\gamma$ -rays. The EBL flux density has been evaluated over almost two decades of wavelengths with a peak amplitude at 1.4  $\mu\text{m}$  of  $\lambda F_{\lambda} = 15 \pm 2_{\text{sys}} \pm 3_{\text{sys}} \text{ nW m}^{-2} \text{sr}^{-1}$ , in between direct measurements and lower limits derived with galaxy counts.

The low energy threshold achieved with the upgrade of the H.E.S.S. array, H.E.S.S. II, will enable the observation of the unabsorbed population of  $\gamma$ -rays and improve the constraints on the intrinsic spectra and thus on the absorption feature. The trough between the COB and the CIB will be characterized by the Cherenkov Telescope Array (CTA, Actis et al. 2011) which will probe energies above 50 TeV. Finally, the increasing size of the sample of blazars detected at very high energies will improve the constraints on the redshift dependence of the EBL and establish a firm observational probe of the thermal history of the Universe.

*Acknowledgements.* The support of the Namibian authorities and of the University of Namibia in facilitating the construction and operation of H.E.S.S. is gratefully acknowledged, as is the support by the German Ministry for Education and Research (BMBF), the Max Planck Society, the German Research Foundation (DFG), the French Ministry for Research, the CNRS-IN2P3 and the Astroparticle Interdisciplinary Programme of the CNRS, the UK Science and Technology Facilities Council (STFC), the IPNP of the Charles University, the Czech Science Foundation, the Polish Ministry of Science and Higher Education, the South African Department of Science and Technology and National Research Foundation, and by the University of Namibia. We appreciate the excellent work of the technical support staff in Berlin, Durham, Hamburg, Heidelberg, Palaiseau, Paris, Saclay, and in Namibia in the construction and operation of the equipment.

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## Appendix A: Study of the systematics

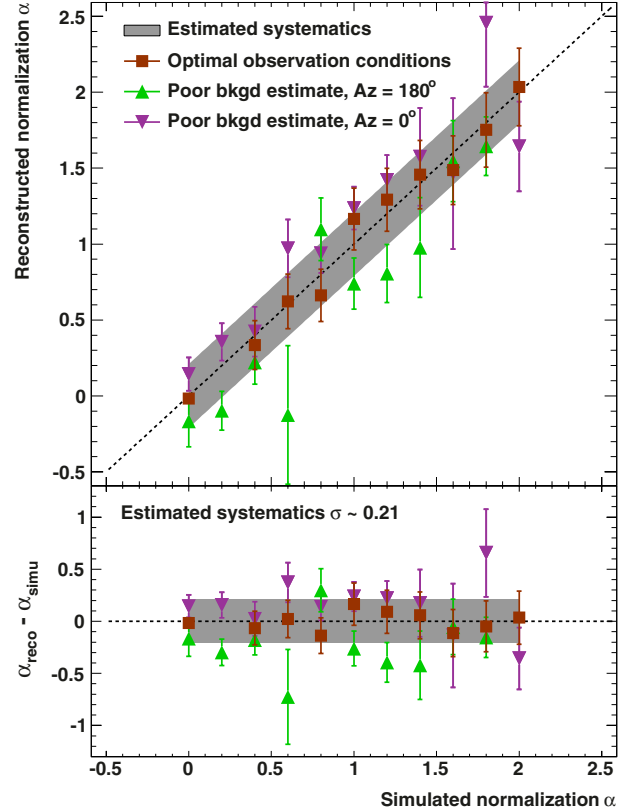
The systematic uncertainties on the EBL measurement with H.E.S.S. brightest blazars are investigated in this appendix. Following [Sinervo \(2003\)](#), two sources of systematics arising from “poorly understood features of the data or analysis technique” (class 2) and two sources of systematics arising “from uncertainties in the underlying theoretical paradigm” (class 3) are identified. The main class 2 systematic uncertainty is evaluated with Monte Carlo simulated air showers imaged by the detector and passing through the whole chain analysis (see, e.g., [Aharonian et al. 2006a](#), and reference therein for a description of the Monte-Carlo simulations). The limited knowledge of the atmospheric conditions is accounted for with a toy model of the detector acceptance and distribution of events. Class 3 systematics are characterized in this study by the choice of template model for the EBL absorption and the selection of the best intrinsic model for each data set. The impact of the latter is evaluated with the data, testing ad hoc intrinsic models, while the former is compared with a concurrent modelling established by [Domínguez et al. \(2011\)](#).

### A.1. Analysis chain

Monte Carlo data (see [Aharonian et al. 2006a](#)) were used to test the analysis chain. Four telescopes triggered events following a PWL of photon index 2 (hardest simulated index) were randomly removed from the simulated data set to create an artificial EBL attenuation. The data set studied was generated for a null azimuth and an off-axis angle of  $0.5^\circ$ . The zenith angle was fixed to  $18^\circ$ , close to the average zenith angle in the H.E.S.S. sky of PKS 2155-304, which is the source with the most important excess of  $\gamma$ -rays in this study (see Sect. 2.2). The EBL optical depth normalization  $\alpha$  was then reconstructed with these samples of events following a spectrum  $\phi(E) \propto E^{-2} \exp(-\alpha \times \tau(E, z))$ , where  $\tau(E, z)$  is the FR08 EBL opacity and  $z$  the redshift of the source, fixed here to  $z = 0.1$  for simplicity.

The background, particularly important for the spectral fit method described in [Piron et al. \(2001\)](#), was fixed to a tenth of the signal – comparable to the value derived for the first data set on PKS 2155-304. The reconstructed EBL normalization  $\alpha$  is shown in the top panel of Fig. A.1 as a function of the simulated EBL normalization. The close match with the identity function strongly supports the reliability of the method employed.

The parameter that seems to affect the analysis chain the most is the background estimation, crucial for the mentioned spectral fit method. Imposing a background equivalent to a fifth of the signal, two samples of simulated events were studied for a null zenith and respective azimuths of  $0^\circ$  and  $180^\circ$ . The azimuth just indexes the data sets, since all azimuth angles are equivalent for a null zenith angle. The corresponding reconstructed EBL normalizations are represented with downward and upward triangles in the top panel of Fig. A.1. The associated error bars represent statistical uncertainties, related to the limited size of the Monte Carlo samples (typically  $10^4$  events), that must be taken into account when estimating the systematic uncertainty. A first (a priori naive) evaluation of this systematic is the average difference  $\alpha_{\text{reco}} - \alpha_{\text{simu}}$  represented in the bottom panel, which reads 0.17 and 0.20 for each sample. A second evaluation is the maximum variation in the measurement  $\Delta$  associated with a Gaussian statistic, yielding one standard deviation systematics  $\Delta/\sqrt{12}$  (see, e.g., [Sinervo 2003](#)) of 0.19 and 0.21, respectively. The estimate chosen is similar to the excess variance estimator developed by [Vaughan et al. \(2003\)](#) for variability. Assuming



**Fig. A.1.** Reconstruction of the EBL normalization with Monte Carlo simulated air showers passing through the analysis chain. Three samples of Monte Carlo events are represented: the first one (orange squares) corresponds to the observation conditions of PKS 2155-304, the second and third (triangles) correspond to a poor background estimation. These two last sets were used to estimate the systematic uncertainty represented with the grey shaded area. *Top panel:* reconstructed EBL normalization as a function of the simulated normalization. *Bottom panel:* residuals, defined as the difference between the reconstructed and simulated optical depth normalizations.

that the rms difference  $D$  between the simulated and reconstructed values is due to both statistical and systematic uncertainties, one would write  $D^2 = \mathcal{V}(\alpha_{\text{reco}} - \alpha_{\text{simu}}) = \langle \sigma_{\text{stat}}^2 \rangle + \sigma_{\text{sys}}^2$ , where  $\mathcal{V}$  is the variance estimator. We thus define the systematic uncertainty estimate as:

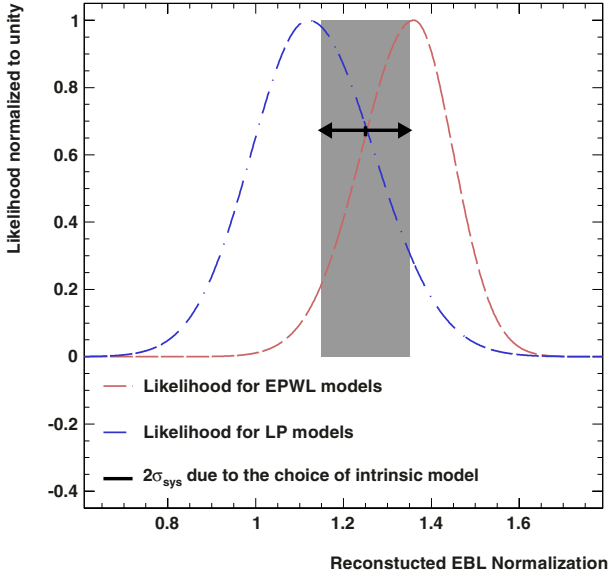
$$\sigma_{\text{sys}} = \sqrt{\mathcal{V}(\alpha_{\text{reco}} - \alpha_{\text{simu}}) - \langle \sigma_{\text{stat}}^2 \rangle}, \quad (\text{A.1})$$

which reads 0.15 and 0.26 for each sample. The global systematic error using both samples,  $\sigma_{\text{sys}} = 0.21$ , is shown in the top and bottom panels of Fig. A.1. This systematics estimate is similar to the two mentioned before, though a bit larger, which suggests a possible slight overestimation.

To ensure that a point-to-point systematic effect does not mimic the EBL absorption as a function of energy, a test was performed with a bright galactic source, the Crab Nebula, and yielded deviations to a null EBL normalization well below the systematic uncertainty derived for the analysis chain.

### A.2. Choice of the intrinsic model

The second systematic uncertainty arises from the choice of the model for the intrinsic spectra. This systematic was assessed on the data by comparing the total likelihood profile derived with a LP for each intrinsic spectrum, on one hand, and with an EPWL,

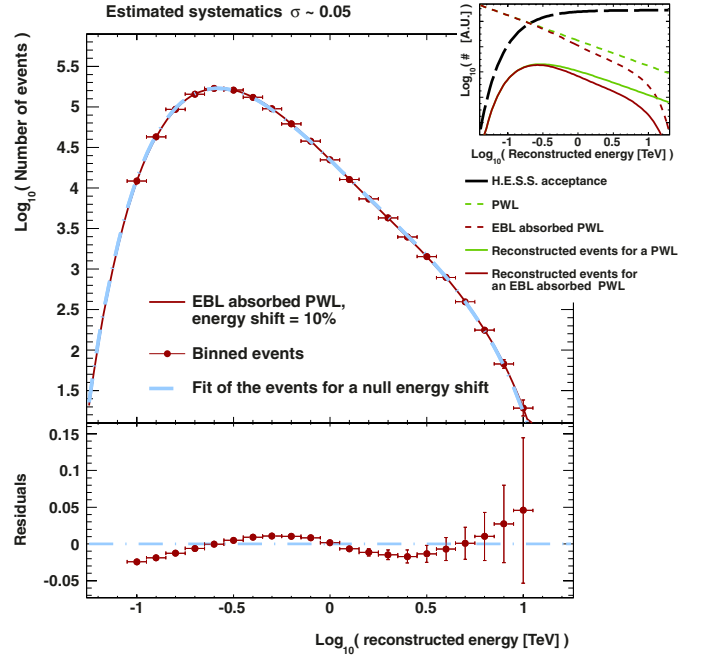


**Fig. A.2.** Likelihood profiles as a function of the normalized EBL opacity. The profiles were normalized to unity for clarity purposes. The dotted dashed curve is derived fitting log-parabolic intrinsic spectra to the data sets, while the dashed curve is derived by fitting exponential cut-off models. The gap between the two profiles due to the intrinsic spectral modelling is represented by the grey shaded area and the double arrow.

on the other. The corresponding likelihoods as a function of the EBL normalization are shown in Fig. A.2, where the maximum was set to unity for clarity. The comparison of third-order models such as ELP and SEPWL would only down the systematic error in the statistical one. The two profiles were fitted with the procedure described in Sect. 3, yielding  $\alpha_{\text{Expcut-off}} = 1.36^{+0.09}_{-0.12}$  and  $\alpha_{\text{LogParabola}} = 1.12^{+0.15}_{-0.13}$ . Using the last systematic estimator described in the Sect. A.1, the difference between these two values due to the statistics is estimated to 0.14 (variance due to uncertainties), and the deviation caused by the systematics is 0.10.

To ensure the reliability of the measurement, three other selection criteria of the intrinsic model were tested. First, the model with the best  $\chi^2$  probability was selected (as in the main method), but the flattest likelihood profile was used in case of ambiguity (e.g. between a LP and an EPWL), yielding a normalization of  $1.18 \pm 0.18$ , preferred at the  $8.9\sigma$  level to a null opacity. A second approach consisted in choosing the simplest model, as long as the next order was not preferred at the  $2\sigma$  level (taking the flattest profile in case of ambiguity), yielding a normalization of  $1.46 \pm 0.11$ , preferred at the  $14.3\sigma$  level to a null opacity. These two criteria do not change the intrinsic model for the data sets on 1ES 0229+200, 1ES 1101-232, Mrk 421 (2), PKS 2005-489 (1 and 2), and PKS 2155-304 (1, 6, and 7). A final test consisted in imposing the most complex model (an ELP) on the other data sets, yielding a normalization of  $1.29 \pm 0.18$ , preferred at the  $7.9\sigma$  level to a null opacity. The above-mentioned systematic uncertainty accounts for the slight changes induced by the selection method and the significance of the result remains almost unchanged.

It is worth noting that the particular attention paid to the intrinsic curvature of the spectra all along the analysis is not superfluous. The likelihood profile obtained assuming that the spectra are described by PWLs is maximum for  $\alpha_{\text{PowerLaw}} = 2.01 \pm 0.07$ . The value derived with such a basic spectral model is significantly above the nominal normalized EBL opacity because intrinsic curvature of the spectra mimics the absorption effect.



**Fig. A.3.** Toy-model of the energy distribution of H.E.S.S. events. The inset in the *top panel* shows the detector acceptance (black line) and the expected distributions of events for a PWL and an EBL absorbed PWL (green and brown lines, respectively). The injected spectra are shifted in energy to model the absorption of Cherenkov light by the atmosphere yielding the distribution of events shown in the *top panel* with brown filled circles. Fitting this distribution with a non shifted model enables the characterization of the atmospheric impact on the EBL normalization estimated to 0.05 for an energy shift of 10%. The residuals of the fit are shown in the *bottom panel*.

### A.3. Energy scale and choice of the EBL model

The atmosphere is the least understood component of an array of Cherenkov telescopes such as H.E.S.S. and can affect the absorption of the Cherenkov light emitted by the air showers. This absorption leads to a decrease in the number of photoelectrons and thus of the reconstructed energy of the primary  $\gamma$ -ray. The typical energy shift, of the order of 10% (Bernlohr 2000), does not affect the slope of a PWL spectrum, which is energy-scale invariant, but impacts its normalization. Indeed, for an initial spectrum  $\phi(E) = \phi_0(E/E_0)^{-\Gamma}$ , an energy shift  $\delta$  yields a measured spectrum  $\phi_{\text{mes}}(E) = \phi_0[(1 + \delta)E/E_0]^{-\Gamma} = \phi'_0(E/E_0)^{-\Gamma}$ , where  $\phi'_0 = (1 + \delta)^{-\Gamma} \phi_0$  is the measured spectral normalization. Since the spectral analysis developed in this study relies on the EBL absorption feature which is not an energy-scale invariant spectral model, the atmosphere absorption impact on the measured EBL normalization is investigated.

A toy-model of the detector and of the atmosphere effect was developed to account for such a systematic effect. The detector acceptance  $\mathcal{A}(E)$  is parametrized as a function that tends to the nominal acceptance value at high energies, as in Eq. (A.2):

$$\log_{10} \mathcal{A}(E) = a \times [1 - b \exp(-c \times \log_{10} E)] \quad (\text{A.2})$$

where  $\mathcal{A}(E)$  is in  $\text{m}^2$ , the energy  $E$  is in TeV, and  $a = 5.19$ ,  $b = 2.32 \times 10^{-2}$ ,  $c = 3.14$  are derived from the fit of the simulated acceptance. The number of events measured in an energy band  $dE$  is then simply  $dN/dE = \mathcal{A}(E) \times \phi(E) \times T_{\text{obs}}$ , where the observation duration  $T_{\text{obs}}$  was fixed to impose a total number of events of  $10^6$ . Typical event distributions for PWL and EBL absorbed PWL spectra are shown in the inset in Fig. A.3. A logarithmic energy binning of  $\Delta \log_{10} E = 0.1$  is adopted and the

uncertainty on the number of events in each energy bin is considered to be Poissonian. To model the effect of the atmosphere on the EBL normalization reconstruction, energy-shifted distributions  $dN/dE = \mathcal{A}(E) \times \phi(E_{\text{shift}}) \times T_{\text{obs}}$  were fitted with a non-shifted model, i.e.  $\propto \mathcal{A}(E) \times \phi(E)$ , with  $E_{\text{shift}} = (1 + \delta) \times E$  and  $\phi(E) \propto E^{-\Gamma} \exp(-\alpha \times \tau(E, z))$ . As mentioned above the effect on the index  $\Gamma$  is null because of the energy-scale invariance, which is not the case for the specific energy dependence of the EBL opacity. A toy-model distribution that was energy shifted is shown in the top panel of Fig. A.3 for a redshift  $z = 0.1$  and an injected EBL normalization  $\alpha = 1$ , corresponding to FR08 EBL modelling. The residuals  $\Delta \log_{10}(N_{\text{events}})$  to the fit of a non-shifted model are shown in the bottom panel.

The reconstructed and injected EBL normalizations differ by less than 0.05 for an energy shift of 10%, while the difference

can go up to 0.11 for an energy shift of 25%. The standard atmospheric conditions required by the data selection motivates the use of the 10% energy shift<sup>4</sup> and thus leads to a systematic error due to Cherenkov light absorption of 0.05.

This toy model of the detector was also employed to compare independent EBL modellings. To probe a reasonable range of models, the lower and upper bounds on the EBL opacity derived by Domínguez et al. (2011) were used for the injected spectrum, while FR08 modelling was fitted to the event distribution. The variation in the reconstructed normalization is estimated to be 0.06 for a redshift  $z = 0.1$ . The small amplitude of the systematic effects of the atmosphere and of the EBL modelling choice (respectively 0.05 and 0.06) justifies a posteriori the use of the simple framework described in this section and does not motivate a deeper investigation.

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<sup>4</sup> Meyer et al. (2010) have even shown that a precision of 5% on the energy scale can be achieved with atmospheric Cherenkov telescopes.